Holographic β function in de Sitter space

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dS entropy and asymptotic freedom

Gibbons Hawking dS_4 is rotated into S_4 The volume of S_4 is $V_{S^4} = 8\pi^2/3H^4$ $S = V_{S^4} 6H^2 / 16\pi G_N = 1/g \quad g = G_N H^2 / \pi$ Distribution function $\rho(\xi(t),\omega) = \sqrt{\frac{4\xi(t)}{\pi q}} \exp\left(-\frac{4\xi(t)}{q}\omega^2\right).$ $S = -\operatorname{tr}(\rho \log \rho) \sim 1/2 \log(1/\xi)$ $<\omega^2>=rac{3g}{4}N(t)=rac{g}{8\xi}$

Renormalization group and FP equation.

 $<\omega^{2}>_{\text{bulk}} = -\frac{3g}{4} \int_{Ha(t)}^{\Lambda} \frac{dk}{k}$ $= \frac{3g}{4} Ht = \frac{3g}{4} N(t).$

Probability density at boundary

IR logs are subtracted

IR log

 $\rho_B = \exp\big(-\frac{3g}{4}\frac{Ht}{2}\frac{\partial^2}{\partial\omega^2}\big)\rho,$

FP equation Renormalization group

$$\dot{\rho} - \frac{3g}{4} \cdot \frac{H}{2} \frac{\partial^2}{\partial \omega^2} \rho = 0,$$

We assume generic Gaussian

$$\frac{\partial}{\partial t}\frac{1}{2}\log(\frac{\xi}{g}) + 3H\xi = 0$$

Master equation

$$\frac{\partial}{\partial N}\log\frac{g}{\xi} = 6\xi.$$

$$\frac{\partial P}{\partial t} = A \Sigma_1^N (\frac{\partial}{\partial x_i})^2 P \qquad P(x_i, 0) = \prod_1^N \delta(x_i - x(0))$$

FP equation g: constant Large N limit

$$\dot{\xi} = -6H\xi^2.$$
 $\xi = \frac{1}{1+6Ht}$

We recall the following identity holds at the horizon exit $t = t_*$ $\dot{\rho}(t_*)e^{\rho(t_*)} = k$

It is nothing but choosing our renormalization scale as $\log k = Ht$.

$$<\omega^2(t)>=\frac{g}{8\xi}=\frac{3g}{4}N(t)$$

$$k^{1-n_s} < \zeta^2 > \propto < \omega^2 > \text{ as } \xi \propto \epsilon.$$

$$g=N^{\frac{m}{2}}, \quad \xi=\frac{m+2}{12N}.$$

Explicit solutions of power potentials V=f(1/n)

$$\epsilon = -\frac{1}{2} \frac{\partial}{\partial N} log(H^2) = \frac{1}{4n\tilde{N}},$$
$$-n_s = \frac{\frac{1}{n} + 2}{2\tilde{N}}.$$

Concave solutions maybe obtained by $m \rightarrow 1/n$: (m,n integers)

$$\epsilon = \frac{m}{4\tilde{N}} = \frac{3m}{m+2}\xi$$

UV fixed point

$$g = \frac{2}{\log N} \left(1 - \frac{1}{\log N}\right), \quad \xi = \frac{1}{6N} \left(1 - \frac{1}{\log N}\right).$$

$$\beta = \frac{\partial}{\partial \log N} g = -\frac{2}{\log^2 N} + \frac{4}{\log(N)^3}$$

$$\beta = 0 \text{ is the UV fixed point}$$

$$\epsilon = -\frac{1}{2} \frac{\partial}{\partial N} \log(g) = -\frac{1}{2gN} \beta_g$$

$$\epsilon \text{ vanishes near UV fixed point}$$









Conclusions and discussions

There may be a pre-inflation era

The entropy grows logarithmically log N N:e-foldings

In the inflationary Universe, the entropy glows power like: (N_e-N)^(-1/n)

The inflationary universe inevitably dominates due to larger entropy

The inflation era lasts until $g \sim 10^{-10} \sim P(\text{density perturbation})$. $N \sim 10^{20}$

The CMB isotropy N \sim 50 is a tiny part of it.

$$\begin{split} & \text{Slow roll and Brownian motion} \\ & \int \sqrt{-g} d^4 x \frac{1}{16\pi G_N} \Big[R - H^2 V(\varphi) - \frac{1}{2} \partial_\mu \varphi \partial^\mu \varphi \Big] \\ & (3H(t)\dot{\varphi} = -\frac{6H^2}{\kappa} \sqrt{\gamma}, \ \varphi \sqrt{\gamma} \kappa = -2\gamma H t, \ 4\epsilon N^2 \sim N \\ & m/4N > \epsilon \sim 1/4N > 1/4nN \\ & \text{Convex potential} \\ & \text{Trans-Planckian problem?} \\ \end{split}$$

$$\delta N \text{ formalism} \qquad \zeta = \delta N = \frac{H}{\dot{\varphi}} \delta \varphi,$$

$$< \delta \varphi(t) \delta \varphi(t') > = \left(\frac{H^2}{4\pi^2}\right) H \delta(t - t').$$

$$< \zeta^2(t) > = < \left(\frac{H}{\dot{\varphi}}\right)^2 \delta \varphi^2 >$$

$$= \frac{1}{2\epsilon M_P^2} < \delta \varphi^2 > = < \frac{H^2}{8\pi^2 \epsilon M_P^2} >$$

$$= \frac{g}{\epsilon} = P$$

